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## Observations of Galaxies with Chandra

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**Abstract.** This talk discusses the growing impact of *Chandra* observations on the study of galaxies. In particular, we address progress in the study of X-ray source populations in diverse galaxies and in the study of the hot ISM.

### 1. Introduction

This talk is a compendium of early *Chandra* (Weisskopf et al 2000) results, published and in progress, and a preliminary view of data available in the public *Chandra* archive. As demonstrated by the large amount of discussion on feedback models in this meeting, and simulations of *Chandra*-like images, the subarcsecond resolution of *Chandra*, combined with its spectral resolution, is going to provide unique answers to the field of galaxy and cluster ‘ecology’ and evolution. *Chandra*’s resolution is optimal for the study of galaxies in X-rays. *Chandra*’s beam is 100 times smaller than that of any past or planned X-ray mission and provides a linear physical size resolution of less than 2 pc for Local Group galaxies and of  $\sim 30$  pc at the Virgo cluster distance. Given the relatively small number and sparse distribution of luminous X-ray sources expected in a normal galaxy (a few hundred at most in the Milky Way and M31, down to X-ray luminosities of  $10^{36}$  ergs/s; Watson 1990; Supper et al 2001), confusion is seldom a problem when observing galaxies in the local Universe with *Chandra*. Moreover, the small beam translates into very efficient source detection, since the background contribution in the beam is virtually nil in most cases: 10 counts make a very significant source. With these capabilities *Chandra* is opening up the field of X-ray population studies in galaxies, and at the same time provides the best facility for in-depth detailed studies of the diffuse hot interstellar emission. These are the two aspects of *Chandra* results on galaxies that I will discuss in this talk.

### 2. X-ray Source Populations

Population studies are often used in astronomy, to both constrain the physical characteristics of a class of sources, and to study their evolution. While statistical studies of the global X-ray properties of galaxies, and of their relation to the overall galaxian emission, could be pursued with pre-*Chandra* capabilities (e.g. see Fabbiano 1989, Eskridge et al 1995, Shapley et al 2001), the study of the properties of the X-ray sources within a galaxy had to be restricted to the

most nearby galaxies. Even so, it was clear that this approach had lots of potential. As I first pointed out in 1995 (Fabbiano 1995), X-ray source population studies are best done outside of the Milky Way: observations of external galaxies provide complete samples of X-ray sources, without the biases in distance and line of sight absorption inherent to Galactic source studies. The first such studies with *Einstein*, followed by similar work with ROSAT, showed that the luminosity functions of X-ray sources may vary in different galaxies, although systematic effects, connected to morphology or star formation activity, could not be discerned at the time. A peculiar result, that was pointed out early on (see Fabbiano 1989) and revisited periodically since, was the presence of ultra-luminous or super-Eddington sources in some galaxies. These are sources with X-ray luminosities well above the Eddington limit for an accreting neutron star ( $\sim 1.3 \times 10^{38}$  ergs/s), suggesting the presence of 10-100s solar mass black holes. Although overall one could not discriminate between single sources and clumps of emission, in a few cases, variability and spectral variability results pointed to compact X-ray binaries (Makishima 1994, Fabbiano 1995).

The first results with *Chandra* have indeed confirmed a variety of functional shapes for the X-ray Luminosity Functions (XLF) of galaxian X-ray sources, and have also uncovered very significant populations of Ultra-Luminous X-ray sources (ULX) in star-forming galaxian regions. A brief summary of individual results follows.

**M33** – The Local Group Sc galaxy M33 has been surveyed with ACIS on *Chandra* (McDowell et al. in preparation). Since this galaxy covers an area of the sky significantly larger than the ACIS field of view, a series of different pointings was needed. So far  $\sim 2/3$  of M33 has been observed with *Chandra*, resulting in  $\sim 120$  sources, with luminosities ranging from  $\sim 10^{39}$  ergs/s (the well-known luminous X-ray nucleus nucleus), down to a threshold luminosity of a few  $\sim 10^{35}$  ergs/s. Excluding the nucleus, all the point-like sources in this galaxy have sub-Eddington luminosities (for a  $1 M_{\odot}$  accretor). Pending more accurate estimates of the completeness at lower luminosities, there is a suggestion of flattening near the detection threshold. Comparison of images in different spectral bands shows a variety of X-ray colors (spectral parameters) for the X-ray sources.

**M81** – The Sb galaxy M81, at a distance of  $\sim 3.5$  Mpc, has overall optical properties very similar to those of M31 (Andromeda), from which it differs chiefly because of its active mini-Seyfert nucleus, that is also the dominant X-ray source. *Einstein* observations, however, suggested that M81 has an intrinsically more luminous XLF than M31 (Fabbiano 1988a). *Chandra* ACIS observations of the central  $8.3' \times 8.3'$  field of this galaxy (Tennant et al 2001) resulted in the detection of 97 point-like sources, down to a limiting luminosity of  $\sim 4 \times 10^{36}$  ergs/s. Tennant et al derive XLFs for the X-ray sources associated with the galactic disk and bulge separately. These two XLFs differ: while the disk XLF follows a power-law over the entire observed range, extending to super-Eddington luminosities, the bulge XLF is steeper, flattening out at luminosities below  $\sim 10^{37}$  ergs/s. These different shapes confirm that bulge and disk X-ray sources belong indeed to different populations, as suggested from Galactic studies (see Watson 1990).

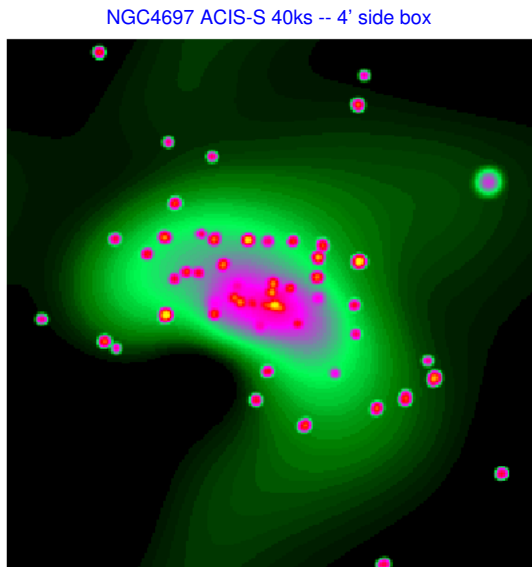


Figure 1. *Chandra* ACIS image of NGC 4697, adaptively smoothed with CIAO ‘csmooth’. Note the population of point-like sources and the ‘warped’ hot ISM.

**E and SO galaxies** – The presence of a bulge-like population of X-ray binaries in early-type galaxies, and their effect on the overall X-ray emission, was first suggested by Trinchieri & Fabbiano (1985; see also e.g. Fabbiano et al 1994, Pellegrini & Fabbiano 1994, for a discussion of their effect on the X-ray luminosity of X-ray faint E and S0s) and has been controversial till now. Although deep ROSAT HRI observations detected some luminous sources in NGC 5128 (Turner et al 1997) and NGC 1399 (Paolillo et al 2001), it is only with *Chandra* that a clear, uncontroversial picture of the X-ray emission of early type galaxies can be obtained. A population of point-like sources is readily visible in the *Chandra* images of NGC 5128 (Centaurus A), where 63 sources were detected in the first HRC image (Kraft et al 2000). Rich populations of point-like sources are also detected in the E NGC 4697 (Sarazin et al 2001; Fig. 1) and S0 NGC 1553 (Blanton et al 2001), and account for  $\sim 1/2 - 2/3$  of the X-ray emission of these X-ray-faint galaxies. The XLFs of NGC 4697 and NGC 1553 both have a break at  $\sim 3 \times 10^{38}$  ergs/s, the Eddington luminosity of a neutron star binary. The *Chandra* image of NGC 1399 reveals a rich population of luminous X-ray sources, prevalently associated with globular clusters (Angelini et al 2001).

**Actively Star-Forming Galaxies** – The XLFs of starburst galaxies tend to be skewed towards higher X-ray luminosities, when compared with those of more aged systems. Examples include NGC 4038/9 (the Antennae; Fabbiano et al 2001; Zezas et al, in preparation), M82 (Zezas et al 2001), NGC 3256 (Lira et al 2001) and more, increasingly being added to the *Chandra* database. Perhaps the most impressive XLF is that of the merging system NGC 4038/9, where 49

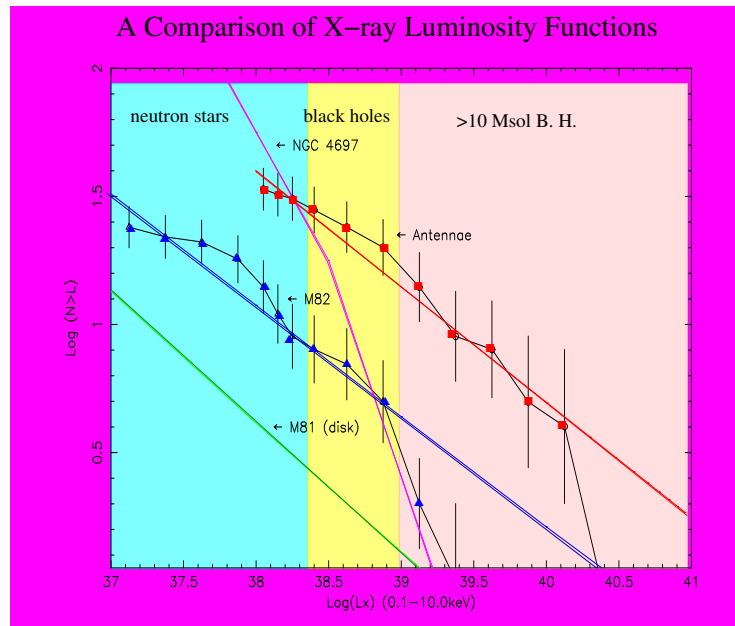


Figure 2. X-ray source luminosity functions in different galaxies.

point-like sources are detected with *Chandra* with luminosities ranging from  $10^{38}$  to  $10^{40}$  ergs/s. There are 14 point-like ULXs sources with  $L_X > 10^{39}$  ergs/s. A number of these sources are also detected in the nuclear region of M82, including the exceptionally luminous, variable source, discussed by Kaaret et al (2001) and Matsumoto et al (2001) and interpreted as a ‘intermediate-mass’ (100s  $M_\odot$ ) black hole binary. The X-ray spectra of the luminous sources in the Antennae are hard and can be fitted with a composite power-law + accretion disk model, reminiscent of the spectra of ultra-luminous sources in more nearby galaxies (e.g. Makishima 2000; Kubota 2001; La Parola et al 2001).

Fig. 2 compares the XLFs of the actively star-forming galaxies M82 and NGC4038/9 with the XLF of the elliptical galaxy NGC 4697 (Sarazin et al 2001) and that of the disk of M81 (Tennant et al 2001). The XLF of NGC 4697 is definitely steeper than the others, and in this resembles that of the bulge of M81 (Tennant et al 2001). This overall difference in shape may be related to the lack of short-lived very luminous sources in older stellar systems. What are these ultra-luminous sources? Although variability and spectra suggest that they are compact binary systems (Zezas et al in preparation), the jury is still out on the nature of the accretor in these binaries: are they ‘intermediate-mass’ black holes (see above), or - more likely- less extreme objects, in beamed X-ray sources (King et al 2001)?

I am sure that in the years to come the study of X-ray populations in galaxies will become an increasingly used tool for understanding properties and evolution of X-ray binaries and their relation to the parent stellar population. Work is already underway on deep images of spiral galaxies (e.g. the work by A. Prestwich and collaborators), and more data are sure to become available in the near future.

### 3. Hot gaseous emission

Hot Inter-Stellar Medium (ISM) can be uniquely detected and studied in X-rays. This hot ISM is found in all kinds of galaxies, and its study may give us information on (i) galaxy structure, e.g. the binding mass of Elliptical galaxies (e.g. Fabricant & Gorenstein 1983); (ii) ecology, by measuring the flux of metal-enriched material from galaxies to the intergalactic medium, via galactic winds; and (iii) evolution, as a way for star formation to affect the energy and entropy balance of clusters and their components (see simulations discussed in this meeting). With *Chandra* we can separate the soft diffuse hot ISM emission from the harder point-source contribution both spatially and spectrally, and we can do detailed spatial/spectral studies of the hot ISM.

#### 3.1. Hot ISM and Superwinds in Starforming Galaxies

Starburst galaxies can be spectacular X-ray emitters, since the enhanced star formation activity results in enhanced X-ray emission and in hot gaseous plumes and galaxian superwinds in the most extreme cases (e.g. Fabbiano 1988b). Examples of this type of galaxies observed with *Chandra* include M82, NGC 4038/9 (Fabbiano et al 2001), NGC 253 (Strickland et al 2000), and NGC 3256 (Lira et al 2001).

In NGC 4038/9, besides the population of bright point-like sources already mentioned, we detect a soft diffuse emission, with a typical optically thin plasma emission spectrum (Fabbiano et al 2001) accounting for about 1/2 of the total X-ray emission ( $\sim 1 \times 10^{41}$  ergs/s). This emission consists of a patchy component associated with the star-forming knots, and a lower surface brightness component extending throughout the system. The extended emission component can be detected farther South than the stellar disks, suggesting a galaxian outflow or superwind.

The angular resolution of the *Chandra* mirrors is such that it is meaningful to pursue direct comparison between *Chandra* and *Hubble* data. A comparison between the ISM component of NGC 4038/9 and the warm ISM visible in  $H\alpha$  (using archival HST WFPC data), suggests a complex and varied multi-phase ISM. While there is a general resemblance of the spatial distribution of the X-ray and  $H\alpha$  emission, we note both regions where the two surface brightnesses closely follow each others, and regions where hot superbubbles fill-in holes in the  $H\alpha$  distribution (Fabbiano et al 2001). The superbubbles we observe in the ISM of the Antennae galaxies are extraordinary, if we compare them with similar features in the ISM of more normal galaxies. Typical X-ray luminosities of bright X-ray superbubbles are in the  $\text{few} \times 10^{39}$  ergs/s range, comparable with the entire thermal emission of the nucleus of the nearby starburst galaxy NGC 253 (Fabbiano 1988b), and  $\sim 20$  times more luminous than 30 Dor.

These observations of nearby active star-forming galaxies provide the detailed data needed to calibrate models of the effect of stellar formation and evolution on the ISM of galaxies. They also provide a local laboratory for the phenomena occurring at the epoch of galaxy formation.

### 3.2. The ‘Feature-rich’ ISM of Early-type Galaxies

The presence of large amounts of ISM in E and S0 galaxies was an X-ray discovery. This ISM is hot and was first convincingly detected with the *Einstein Observatory*. A short history of the observations and modelling of this ISM, and the ensuing controversies can be found in Fabbiano & Kessler (2001). With *Chandra*, we are obtaining a new detailed look that reveals a wealth of structure in what were by-and-large considered fairly uniform hot halos shaped by the galaxy’s gravitational field. Examples of ‘disturbed’ hot halos are provided by the *Chandra* observations of NGC 4636, NGC 5044 and NGC 4697.

**NGC 4636** – This Virgo elliptical is one of the galaxies with largest, most extended X-ray halos. The *Chandra* observation (proposed by R. Mushotzky; data are now in the public *Chandra* archive), reveals both a rich population of individual galactic X-ray sources, including a ring-like concentration in the nuclear region, and a strongly defined, 8 kpc long, spiral-like feature in the X-ray halo (Fig. 3). Recently, Jones et al (2001) propose that this structure may be due to shocks driven by a past nuclear outburst. Ciotti and Ostriker (1997, 2001) developed a model in which recurrent outburst and cooling flow accretion episodes may occur in early-type galaxies, giving rise to AGN/quiescent cycles.

**NGC 5044** – This elliptical galaxy, observed by P. Goudfrooij, and also in the *Chandra* public archive, provides another example of a large, centrally disturbed, X-ray halo (Fig. 4). In this case, a central radio source is present, and the X-ray feature could be related to it. A disturbed inner halo is also present in NGC 1399, another early-type galaxy with nuclear radio emission, where some of the features suggest interaction between radio lobes and hot ISM (Paolillo et al 2000).

**NGC 4697** – This is the X-ray faint galaxy whose X-ray source population was discussed earlier in this paper (Fig. 1). As reported by Sarazin et al (2001), the *Chandra* image also shows a soft extended component of the X-ray emission. This hot ISM does not follow the general symmetry of the stellar distribution, but it is both more extended in its radial behaviour and warped, suggesting some dynamical interaction.

Clearly these are only a few first results. It is not difficult to predict, based on these data, that *Chandra* observations of the gaseous component of E and S0 galaxies will provide novel insight on their evolution and physical properties.

## 4. Conclusions

It is certainly way too soon to evaluate the effect of high spatial/spectral resolution *Chandra* X-ray observations in our understanding of galaxies and their environment, since only a very small amount of data have been fully analyzed and more observations are continually being done. However, a first look shows that *Chandra* has more than fulfilled its promises.

High resolution spectrally-resolved images have solved many outstanding issues, in a very direct and simple way. Most of these issues were connected with the interpretation of lower resolution data of E and S0 galaxies. I am pleased to say that the conclusion, put forward by me and my collaborators,